

Multiwavelength Observations of one Galaxy in Marano Field

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Abstract We report the multi-wavelength observations of one intermediate redshift ($z=0.3884$) galaxy in the Marano Field. These data include ISOCAM middle infrared, VLT/FORS2 spectroscopic and photometric data, associated with the ATCA 1.4 GHz radio and ROSAT PSPC X-ray observations from literature. The Spectral Energy Distribution obtained by VLT spectroscopy exhibits its early-type galaxy property, while, in the same time, it has obvious [O III]5007 emission line. The diagnostic diagram from the optical emission line ratios shows its Seyfert galaxy property. Its infrared-radio relation follows the correlation of sources detected at $15\ \mu\text{m}$ and radio. It has a high X-ray luminosity of $1.26 \times 10^{43}\ \text{erg s}^{-1}$, which is much higher than the general elliptical galaxies with the similar B band luminosity, and is about 2 orders of magnitude higher than the derived value from the star forming tracer, the FIR luminosity. This means that the X-ray sources of this galaxy are not stellar components, but the AGN is the dominant component.

Key words: Galaxies: active — Galaxies: elliptical and lenticular, cD — Galaxies: Seyfert Galaxies: starburst

1 INTRODUCTION

The Marano Field is a ~ 1 sq. deg. region in the southern sky around $\alpha(2000)=03^{\text{h}}15^{\text{m}}09^{\text{s}}$ and $\delta(2000)=-55^{\circ}13'57''$. A series of papers have presented the multiwavelength data of this field. The first deep survey in this field was an optical search for quasars, selected by colour, grism spectra, and optical variability. 64 quasars and AGN were detected at the limiting magnitude of $B_J=22$ in a survey area of $0.69\ \text{deg}^2$ (Marano, Zamorani & Ziteklli 1988; Zitelli et al. 1992).

In the soft X-ray band (0.5–2 keV), the field has been observed for ~ 56 ksec with the Positional Sensitive Proportional Counter (PSPC) on ROSAT. Zamorani et al. (1999) presented such X-ray data and the optical identifications (the ESO NTT for a set of U, B, V, R, J, F bands). In the inner $15'$ field, they detected 50 X-ray sources with $S_x > 3.7 \times 10^{-15}\ \text{erg cm}^{-2}\ \text{s}^{-1}$. Among them, 33 are AGNs and two are radio galaxies.

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Deep radio surveys at 1.4 GHz and 2.4 GHz were carried out in Marano Field with the Australia Telescope Compact Array (ATCA). At the limiting flux of ~ 0.2 mJy, 63 sources were found at 1.4 GHz and 48 sources at 2.4 GHz (Grupponi et al. 1997).

However, the overlap of the radio sources with the ROSAT X-ray sources is relatively small. A cross-correlation analysis has been presented by Grupponi et al. (1996). Three of the five detected X-ray sources have been spectroscopically identified: two of them are quasars and the remaining is an early type galaxy. This early type galaxy is the object we are studying here.

XMM-Newton survey was also carried in an exposure time ~ 80 ksec in the Marano Field (Girdke et al. 2003). 453 new X-ray objects were discovered except the 50 X-ray objects detected by ROSAT. The optical identifications for XMM-Newton survey have also been done (Lamer et al. 2003). However, there are no public data available for the individual object in the both observations at this moment.

We have performed the VLT/FORS2 spectroscopic observation on some objects in the Marano Field. The high quality VLT spectra must provide important information on the intrinsic properties of the objects. Also, deep ISOCAM MIR Guaranteed Time Extragalactic Survey has been done in a 90 arcmin² region. The source counts presented by Elbaz et al. (1999) have shown strong evolution of the cosmic infrared background.

In this paper, we study the properties of one $15 \mu\text{m}$ -selected galaxy, UDSR09, in this field. We present its high quality VLT/FORS2 optical spectrum and ISOCAM MIR data, associated with the X-ray and radio data from literature. The meaning of UDSR09 marks the ninth slit of the FORS2 spectroscopic observation in the Ultra-Deep-Survey-ROSAT field (Liang et al. 2004). This object is the only one cross-correlated from the multi-wavelength data mentioned above. We discuss its radio-IR correlation, the X-ray luminosity L_X versus B band luminosity L_B , L_X versus FIR luminosity L_{FIR} as well. This paper is organized as follows, The multiwavelength data of the galaxy are presented in Sect.2, the results and discussions are given in Sect.3. Sect.4 gives the conclusion. Throughout this paper, a cosmological model with $H_0=70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M=0.3$ and $\Omega_\Lambda = 0.7$ has been adopted.

2 MULTIWAVELENGTH DATA

2.1 IR data from ISOCAM

The very deep ISOCAM follow-up in the Marano field has been done reaching a flux limit of about $50 \mu\text{Jy}$. In the 90 arcmin² field, 75 objects have been detected (Elbaz et al. 1999; Elbaz et al. 2005, in preparation), 35 of them have been observed by using VLT/FORS2 spectroscopy (Liang et al. 2004). Most of the ISOCAM galaxies are star forming galaxies ($>77\%$), rather than AGN (Liang et al.2004; Fadda et al. 2002; Elbaz et al. 2002; Franceschini et al. 2003). The IR data of the sample galaxy can be found in Table 1.

Table 1 The IR Data of the Sample Galaxy

Object	RA	DEC	z	$f(S_{15})$	$\log(\frac{L_{\text{IR}}}{L_\odot})$	SFR_{IR}
	(2000)	(2000)		(μJy)		($M_\odot \text{ yr}^{-1}$)
UDSR09	3 14 56.1	-55 20 08	0.3884	608.9	11.30 ± 0.08	34.21

2.2 Optical data from VLT

Spectrophotometric observations of the ISOCAM-detected galaxies in the Marano Field were obtained by using the ESO 8m VLT using the FORS2 with R600, I600 at a resolution of 5 \AA and

covering the possible wavelength range between 5000 and 9000Å. The slit width was 1.2'' and the slit length was 10''. Spectra were extracted and wavelength-calibrated using the IRAF package. Flux calibrations were done using 15 minute exposures of photometric standard stars (Liang et al. 2004). We study the properties of the galaxy, UDSR09, the object superimposed on the ninth slit in the VLT observation run.

Using the GIM2D software package (Simard et al. 2002), we analyze the optical image of UDSR09 obtained from VLT/FORS2 and try to obtain some morphology parameters, including the bulge fraction, bulge effective radius, disk scale length and inclination i (the disk axis to the line of sight) (see Table 2).

Table 2 The Optical Data: Emission Line Ratios and Morphology Parameters

Object	M_B	A_V	$\log(\frac{[OII]}{H\beta})$	$\log(\frac{[OIII]}{H\beta})$	b	Re	Rd	i
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
UDSR09	-21.58	3.99	1.61±0.09	1.35±0.03	0.77±0.02	5.61±0.30	3.34±0.30	57±5

Notes: b refers bulge fraction; Re, bulge effective radius; Rd, exponential disk scale length; i , disk inclination.

2.3 Radio data from ATCA from reference

Gruppioni et al. (1997) presented the radio observations of the Marano Field carried with the ATCA at 1.4 and 2.4 GHz with a limiting flux of ~ 0.2 mJy at each frequency, extracted from an area of ~ 0.36 deg². 63 objects were detected at 1.4 GHz. One of the galaxies with multiwavelength data available is studied in this paper (see Table 3, Cols.(2)–(6)).

Table 3 The Radio and X-ray Data from Literature

Object	Radio ^a					X-ray ^b		
	object	$S_{\text{Peak},1.4}$ (mJy)	$S_{\text{Total},1.4}$ (mJy)	$S_{\text{Peak},2.4}$ (mJy)	$S_{\text{Total},2.4}$ (mJy)	object	Net counts	S_x (10^{-14} erg cm ⁻² s ⁻¹)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
UDSR09	38B	0.36±0.03	0.57±0.08	0.43±0.04	0.44±0.12	X021-5	102.3	2.36±0.25

Notes: “a” is for (Gruppioni et al. 1997) and “b” is for ROSAT PSPC from Zamorani et al. (1999). Cols.(2) and (7) are the corresponding names in the two references.

2.4 X-ray data from ROSAT from reference

The ROSAT PSPC observation in the Marano Field detected about 50 X-ray sources in the inner 15 arcmin radius circle at a flux limit of $\sim 3.7 \times 10^{-15}$ erg cm⁻² s⁻¹. 33 of them are AGNs and two are radio galaxies (Zamorani et al. 1999). One of the radio galaxies is the target source in this paper (see Table 3, Cols.(7)–(9)).

3 RESULTS AND DISCUSSION

The derived results and characteristic parameters for the target galaxy are presented in Table 1–4. Col.(1) of all the tables gives the name of the galaxy, UDSR09, following Liang et al. (2004).

3.1 IR data

Table 1 illustrates the basic data of the galaxy as follows, RA and DEC in 2000 epoch, redshift, ISOCAM $15\ \mu\text{m}$ flux, far-infrared luminosity and Star Formation Rate (SFR) derived from IR luminosity. This galaxy is a luminous infrared galaxy in the $z \sim 0.4$ intermediate redshift Universe, with a high SFR of $\sim 34 M_{\odot} \text{ yr}^{-1}$.

3.2 Optical data

Figure 1a shows the optical spectral energy distribution of the galaxy obtained from VLT/FORS2. It shows an elliptical galaxy property, and has strong metal absorption lines, for example, Ca H K, G-band, Mg H, Na D lines, but weak Balmer absorption lines. In particular, it shows obvious [OIII]5007 emission line, as well as [OII]3727 emission line. The $H\beta$ emission line is weak.

In Table 2, Col.(2) gives the B band absolute magnitude in the Vega system, Col.(3) gives the dust extinction by considering energy balance between IR and $H\beta$ luminosities, Cols.(4) and (5) are the emission line ratios, which exhibit its Seyfert 2 galaxy property, which is shown in the diagnostic diagram of Fig. 1b (also see Liang et al. 2004).

We use GIM2D (an exponential disk and a de Vaucouleurs $R^{\frac{1}{4}}$ law for bulge) to try to obtain the morphology parameters of the galaxy (Cols.(6)–(9) in Table 2). The interesting thing is the bulge effective radius R_e is larger than the disk scale length R_d , which means the morphological parameters of this galaxy cannot be decomposed into simple bulge/disk decompositions.

3.3 Radio-IR correlation

Table 4 (Cols.(7),(8)) illustrates the $15\ \mu\text{m}$ middle infrared and 1.4 GHz radio luminosities of the galaxy, which generally follow the radio-IR correlations presented by Elbaz et al. (2002) (their fig. 9a) and Gruppioni et al. (2003) (their fig. 1) for $15\ \mu\text{m}$ -detected radio galaxies. Our Fig. 1c shows such correlation of it. And the corresponding q value (from Eq.(2) of Elbaz et al. 2002) is 1.74, a little lower than those of the local galaxies (2.34) and the ISOCAM intermediate- z galaxies (2.3) shown in fig. 9b of Elbaz et al. (2002).

Table 4 Some Derived Luminosities and Parameters

Object	$\log(\frac{L_B}{L_{\odot}})$	flux([OIII]) ($\text{erg cm}^{-2} \text{ s}^{-1}$)	$L_{[\text{O III}]}$ (erg s^{-1})	L_X (erg s^{-1})	L_{IR} (erg s^{-1})	$\nu L_{\nu}[15\ \mu\text{m}]$ (L_{\odot})	$\nu L_{\nu}[1.4\text{GHz}]$ (L_{\odot})	q
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
UDSR09	9.98	9.34×10^{-17}	4.9×10^{40}	1.26×10^{43}	7.63×10^{44}	2.13×10^{10}	6.95×10^5	1.74

3.4 X-ray luminosity versus B band luminosity

If the X-ray emission sources in the early-type galaxies are stellar in origin, the total X-ray luminosity from these sources should scale with B band luminosity L_B . While emission from other sources, such as hot gas, may not be so directly linked to stellar population (Irwin & Sarazin 1998; O’Sullivan et al. 2001).

Table 4 gives the derived L_X (0.5–2 keV from ROSAT) and L_B of the target galaxy. The galaxy has high L_X , $\sim 1.26 \times 10^{43} \text{ erg s}^{-1}$, which is much higher than that of the general early-type galaxies with similar blue band luminosity (see our Fig. 1d). We consider that, if all blue stars were embedded in the ionized gas, the L_B may be underestimated by 3 mag (Liang et al.

2004). This will lead $\log L_B$ to increase ~ 1 dex, but still correspond to much lower L_X than the observed one at the similar L_B .

Since the target galaxy shows much higher L_X than the general early-type galaxies at the given L_B , the source of its X-ray emission may not be the stellar components. There must be other sources causing the unusually high X-ray luminosity, e.g. hot gas or the central buried Active Galactic Nuclei (AGN). There may exist a Black Hole at its center.

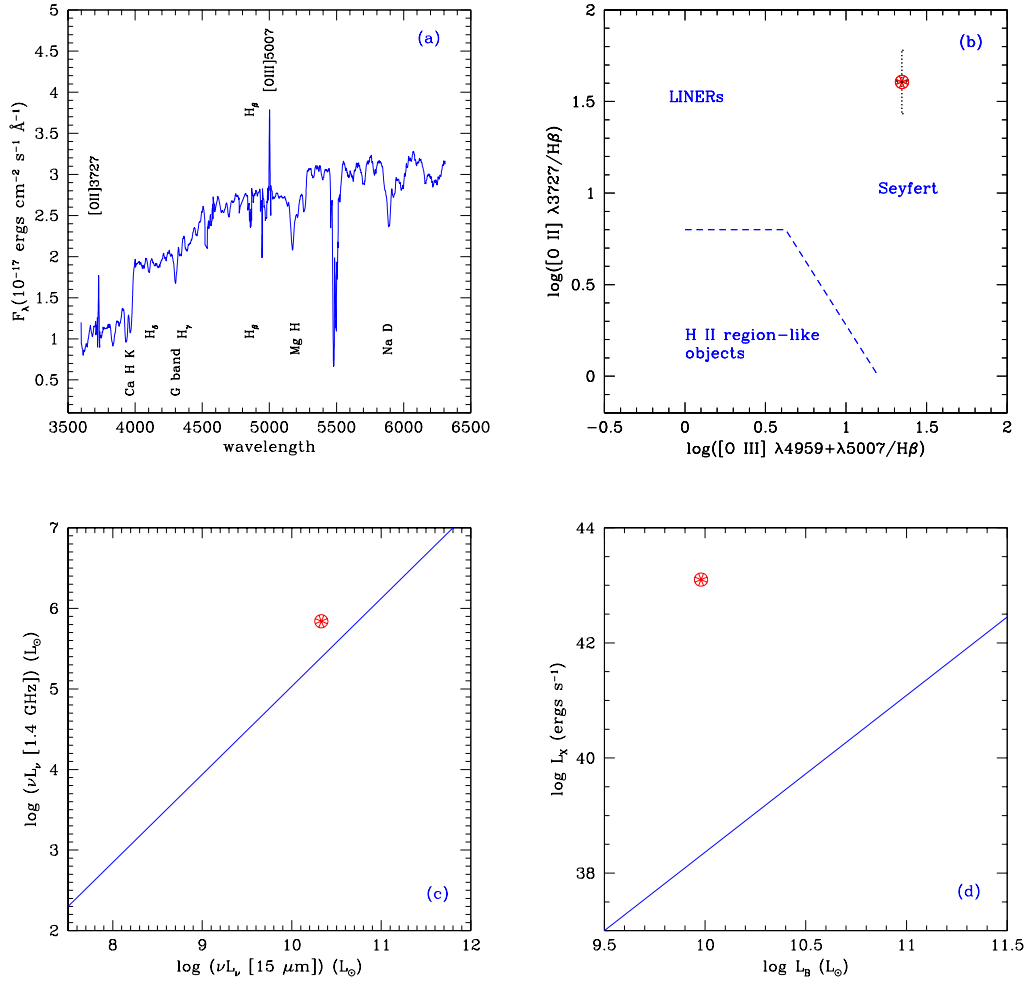


Fig. 1 (a) The VLT/FORS2 optical spectrum of the galaxy (the top left); (b) The diagnostic diagram from optical emission line ratios (the top right); (c) The radio-IR correlation (the bottom left), and the solid line is from Gruppioni et al. (2001) for the best-fitting obtained for the radio sources detected by 15 μm in the ELATS southern fields; (d) The X-ray luminosity (0.5–2 keV) versus the blue optical luminosity (the bottom right), and the solid line is taken from Fig. 1 of Irwin & Sarazin (1998) for the best-fitting power-law relation for early-type galaxies.

3.5 X-ray luminosity versus FIR luminosity: the SFR

To further understand the sources of X-ray emission, one way is to compare the L_{FIR} with L_{X} , hence the derived SFRs. Ranali et al. (2003) studied the X-ray luminosity as a SFR indicator for the local star forming galaxies: $\log(L_{0.5-2\text{keV}}) = (0.87 \pm 0.08)\log(L_{\text{FIR}}) + 2.0(\pm 3.7)$.

The L_{FIR} of the galaxy is $\sim 4.0 \times 10^{44} \text{ erg s}^{-1}$ (from L_{IR} by using Eq.(4) of Elbaz et al. 2002). This corresponds to $L_{\text{X}}(0.5 - 2 \text{ keV}) \sim 6.3 \times 10^{40} \text{ erg s}^{-1}$ as suggested above. This is about two orders of magnitude lower than the L_{X} derived directly from ROSAT flux ($1.26 \times 10^{43} \text{ erg s}^{-1}$). This obvious discrepancy confirms that the X-ray emission of the sample galaxy may not be due to star forming, there must be other sources responsible for it, e.g. the buried AGN.

4 CONCLUSIONS

We present the multiwavelength data of one galaxy in Marano Field, including the ISOCAM middle infrared, VLT spectroscopic and photometric data obtained by us, associated with the ATAC radio 1.4 GHz and ROSAT PSPC X-ray data (0.5–2 keV) from the literature.

The high quality VLT spectrum show the obvious [OIII] emission line. The [OIII], [OII] and $H\beta$ emission lines diagnostic diagram shows its Seyfert 2 galaxy property. Its IR and radio data relation follows the radio-IR correlations of local and ISOCAM intermediate- z galaxies.

However, its X-ray luminosity is much higher than that of general early-type galaxies at the given B band luminosity, and much higher than the star forming tracer, FIR luminosities, which means that the sources of its X-ray emission should not come from stellar components, and the AGN is the dominate component.

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